Keck Adaptive Optics Note 551
(DRAFT)
Keck Next Generation Adaptive Optics
WFS sub-system conceptual study report

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1. Introduction

The Next Generation Keck Adaptive Optics system is a multiple guide star Adaptive Optics system with a two stage reflective OAP relay designed to work at wavefront errors as low as 90 nm. The system is envisioned to have four kinds of wavefront sensors, namely, one NGS WFS, 9 LGS WFS's, 3 TT(FA)s (2TTs and 1 TTFA), and two truth sensors. This is a report of a conceptual study undertaken. Though throughout the study both the *1-tier* and the *2-tier* optical designs were borne in mind while developing concepts, this report deals with the *down selected 1-tier* optical design alone.

2. Work scope definition:

3.2.3.5 Overall WBS Dictionary definition: Develop a design concept for each of the required NGAO wavefront sensors

3.2.3.5.1 High Order LGS Wavefront Sensors: Given the functional and performance requirements, develop a design concept for the laser guide star high order wavefront sensors. ~~Take into consideration the possible need for both open and closed loop wavefront sensing~~.

3.2.3.5.2 High Order NGS Wavefront Sensor: Given the functional and performance requirements, develop a design concept for the natural guide star high order wavefront sensor(s~~). Take into consideration the possible need for both open and closed loop wavefront sensing~~. Include consideration of ADC packaging (ADC design is covered in WBS 3.2.3.8).

3.2.3.5.3 Low Order NGS Wavefront Sensors: Given the functional and performance requirements, develop a design concept for the low order natural guide star wavefront sensors for the purpose of determining tip/tilt and other low order modes in laser guide star observing mode. ~~Take into consideration the possible need for both open and closed loop wavefront sensing.~~ Include consideration of ADC packaging (ADC design is covered in WBS 3.2.3.8).

3.2.3.5.4 Calibration Wavefront Sensor *(Called truth sensor in this document)*: Given the functional and performance requirements, develop a design concept for the calibration wavefront sensor which will use natural guide star light as a truth wavefront. This sensor will be periodically used to reset the references of the high order wavefront sensors in laser guide star mode. Include consideration of ADC packaging (ADC design is covered in WBS 3.2.3.8).

Inputs to the study:

1. Optical design of the Cascaded relay.
2. FRD
    - Type of each WFS (SH/ PYR)
    - What order/ # of sub-apertures.
    - The position each sensor in the Optical Relay.

        - The FoR for each sensor
        - Positioning accuracy
        - Choice of detector(s) for each WFS (pixel size).
        - Pixel geometry specifics like guard bands, pixel geometry and spot size (for LGS with appropriate elongation) for each WFS
        - Centroiding accuracy, dynamic range, and linearity specifications from FRD for each WFS
        - The Field Stop/ Spatial filter specification
        - TT sensor specification (FoV, dynamic range, etc.)
    3. SRD (specifically input of the type of sources on which tip-tilt sensor needs to work and performance margin for binaries, elongated/ asymmetric sources.
    4. NGAO System Architecture Definition (KAON 499)
    5. Mechanical drawing(s) w/ space constraints and packaging issues clearly stated for the of Cascaded relay.
    6. Specification on pick-offs for the WFSs (including the ones shared by the TT(FA) sensors inside the d-NIRI) and rotation if necessary. (input must come from 3.2.3.11)
    7. Wavefront sensor error budget spreadsheet.

3. Products:

1. Conceptual optical design(s).
2. Feed into relevant sections of FRD version 2.0 (in particular update TT sensor requirements and performance based on the type of source).
3. LGS pick off mechanism concepts.
4. Conceptual designs and first order optical design for the LGS WFSs, TT(FA) sensors.
5. First order Mechanical packaging.
6. Preliminary mechanical design and 3D model (at least a cartoon showing the envelopes occupied by the WFSs).
7. Acceptance and completeness of concepts and conceptual design with information on what needs to be done during the preliminary design phase.
8. Update the terms in the error budget spreadsheet based on conceptual design.
9. Documentation for all the above.

4. Methodology:

1. Liaise with opto-mechanical team to understand the optical and mechanical constraints.
2. Based on the inputs from section 2, a first order optical design shall be worked out and shared with the rest of the WFS team for scrutiny. This will be documented and sent to the EC for further inputs.
3. Conceptual designs for LGS pick offs will be worked out based on work done by the IWG, other projects like Gemini MCAO, and other MOS pick off options and

Palomar Tomograph. Mechanism(s) for registering each LGS sensor to the DM will be conceptualized.

4. One day meeting to understand risks and look at the acceptance of the work and design with the WFS design team.
5. Documentation of the designs with design risks stated.
6. Update the FRD and Error Budget Spreadsheet as and when they need updating based on how the conceptual design evolves.

3. Input parameters

WFS type	Location	Sensing wavelength (nm)	Input PS (um/")	# of sub-apertures	Detector PS("/pixel)	Filters	Comments
NGS HOWFS	near NF sci. inst.	400-900	2254	32x32, 64x64	1.5	no filter	Steering mirrors for OSM
LGS HOWFS	after WF relay	589 nm	727	16x16, 32x32, 64x64	1.45	no filter	Has to track Na layer, tilted focal plane, field dependent aberrations. Steering mirrors for OSM
TT	inside d-NIRI	1.16 - 1.33	727	1x1	0.030	J, H and J & H	will use d-NIRI OSM
TTFA	inside d-NIRI	1.16 - 1.33	727	2x2	0.030	J, H and J & H	will use d-NIRI OSM
Truth sensor	1 sensor after the NF and 1 after the WF	400-900(NF) 1000-2400 nm (WF)	2254 and 727	5x5	.65 and .200	no filter	will use field steering and d-NIRI pick offs

Table 1: Baseline design parameters for the NGAO WFSs. The NGS HOWFS, the LGS HOWFS, the TTFA and the truth sensor are all SH sensors. The TT sensor works like a single quad-cell

WFS type	FoR	Baseline detector choice	Pixels/sub-ap	Guard bands?	spot size delivered to the sensor	linearity	field stop	Comments
NGS HOWFS	50"	Visible CCID56	4	none at 64x64		spot size ~ pixel size	specs TBD	Assumes Peltier package
LGS HOWFS	174"	Visible CCID56	4	none at 64x64			specs TBD	Assumes Peltier package

TT	150"	IR detector with selectable pixels and RG capability	4-many	NA				many pixels for extended objects
TTFA	150"	IR detector with selectable pixels and RG capability	4-many	at least a few pix.			specs TBD	2x2 lenslet
Truth sensor	150" (50")	IR detector with selectable pixels and RG capability (visible detector with low dark current)	4	few pixels			specs TBD	

Table 2 Other wavefront sensor parameters of interest

4. The LGS wavefront sensor:

There are a total of 9 LGS wavefront sensors in the baseline NGAO design. The asterism is novel in the sense that there is a fixed central guide star with a 5 star asterism around it on the vertices of a regular pentagon and There are 3 additional point and shoot lasers that can be pointed anywhere in the FoR to MOAO sharpen the 3 TT(FA) stars (c.f. Fig. 1). The radius of the asterism that conforms to a regular pentagon can be varied from 20-174". It is from these

A lenslet switching mechanism as shown in Figure 6 will be used to support multiple pupil sampling scales.

The LGSF team has come up with a scheme to keep the lasers fixed on the sky (c.f. Figures 6 and 7). The LGS pick offs use a theta-phi rotating mechanism that keeps the path length constant at all times. The schematic for this mechanism is shown in Figures 3, 4 and 5.

The motion control required for a single LGS channel is shown in Fig.2. The overall list of mechanisms is documented by Erik Johansson as part of the non-time real time supervisory control documentation based on our input. The list of motion control required by the LGS WFS is given in Table 3 with travel and tolerances.

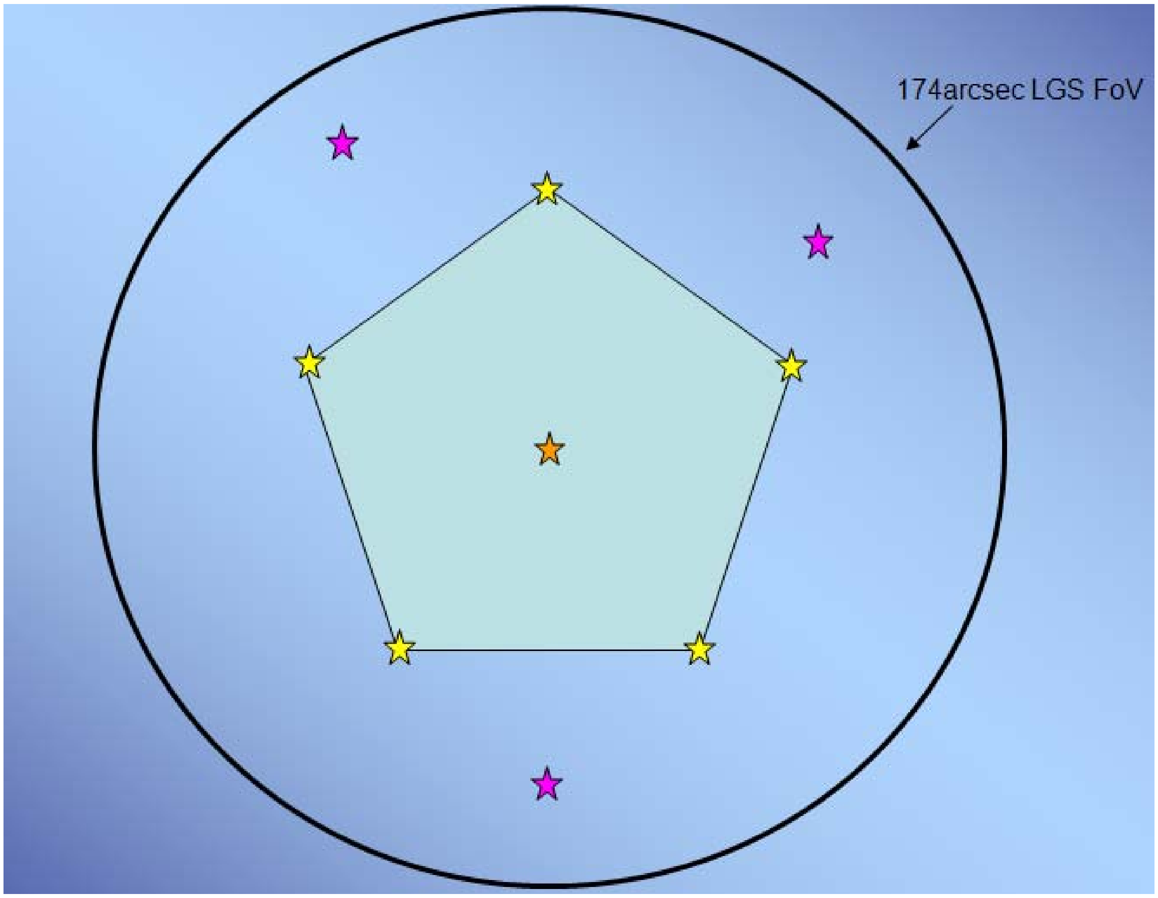


Figure 1 LGS guide star asterism showing fixed central guide star, 5 equi-angle LGS star with radiuses varying from 10-150" and 3 roaming guide stars with 174" FoV.

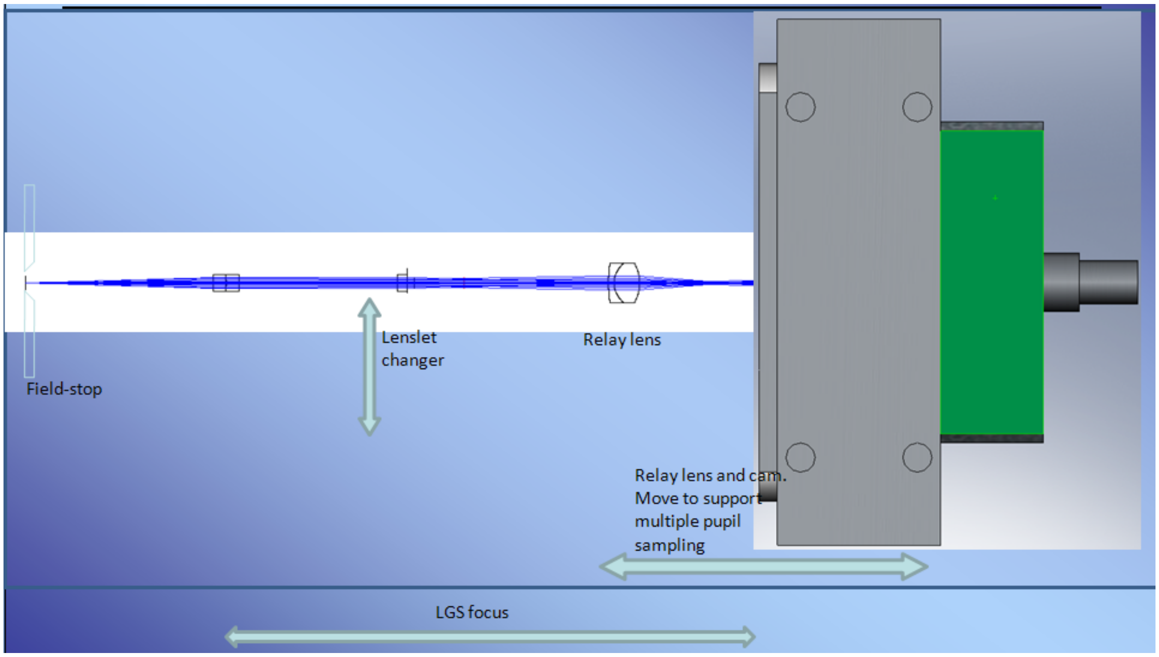


Figure 2 Schematic of a single LGS WFS channel on which 1^{st} order design was based on.

Location	Type	Range of travel	Min. step size (est.)
LGS channel (individual unit focus)	delta Z	~10 mm	10 um
LGS lenslet changer	deltaX, Y	delta X = 20 mm, delta Y = 2 mm	1 um
LGS channel relay and camera focus for pupil sampling change	delta Z	5 mm	1 um
Pick off (theta mech)	delta theta	360 deg.	0.2 (UP)/ 115" (LA)
Pick off (phi mech.)	delta phi	360 deg.	0.2" (UP)/ 300" (LA)
LGS WFS unit overall travel	delta Z	130 mm	1 mm
Rotation of the WFS (sans the central WFS)	delta (theta-phi)	360 deg.	100"

Table 3 Motion control parameters for the LGS WFS.

Pick-off mechanism:

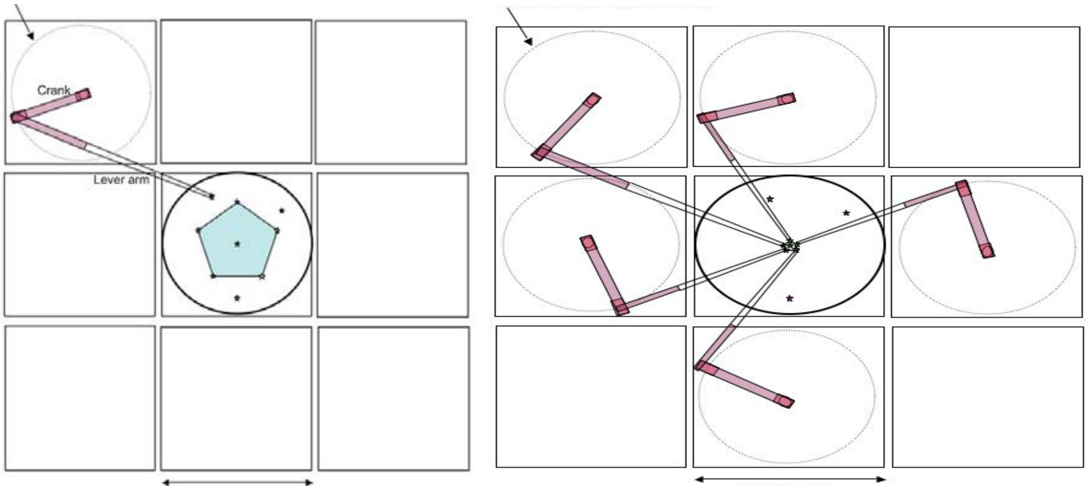


Figure 3 Plan view of the pick-off scheme showing the theta-phi mechanism. The 8 sensors that are around the central star will need to have a rotation stage that ensures that the lenslet array, the DM and the corresponding sub-aperture pixels are all aligned to one-another at all times. Though this is not an active control, every time a new LGS star is picked off, we will need to reorient the WFS in rotation to maintain the one-to-one mapping of the DM actuators to lenslets.

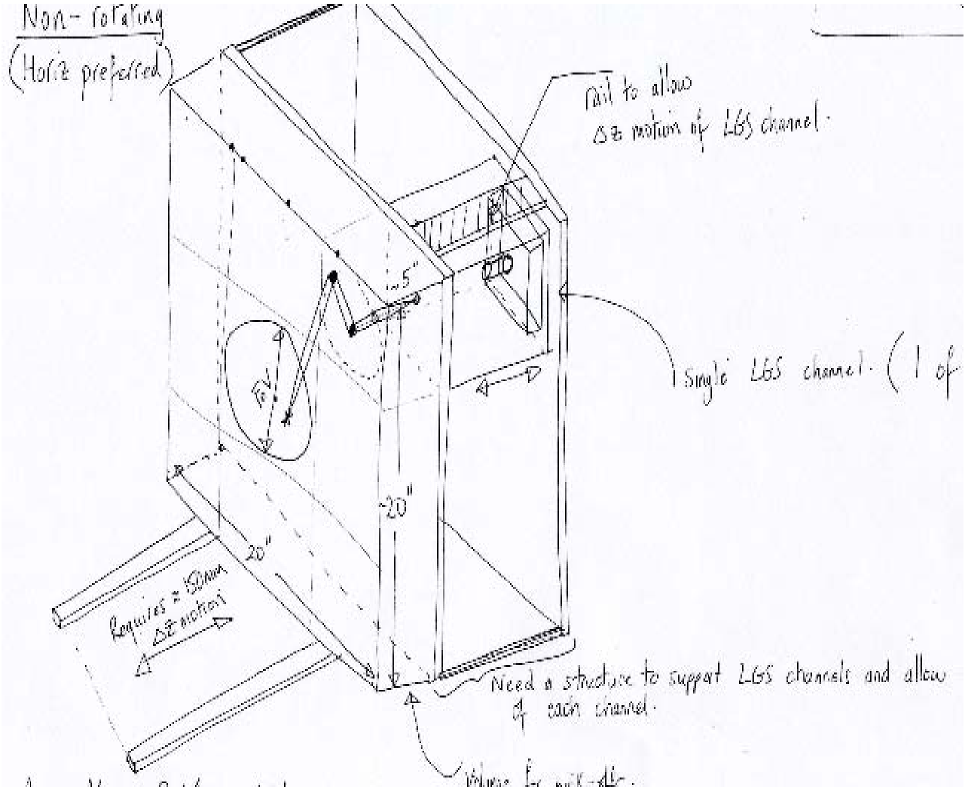


Figure 4 A architect’s view of the LGS WFS unit showing 1 out of 9 channels with a single pick-off arm.

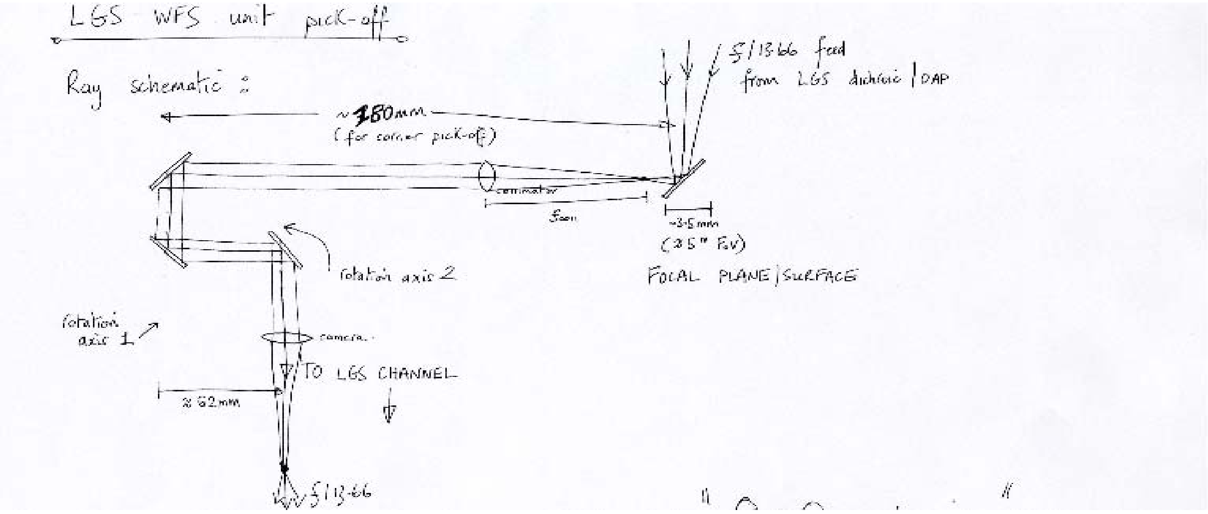


Figure 5 Schematic that unravels the theta-phi mechanism.

Why does the whole LGS WFS unit not need to actively rotate during an observation?

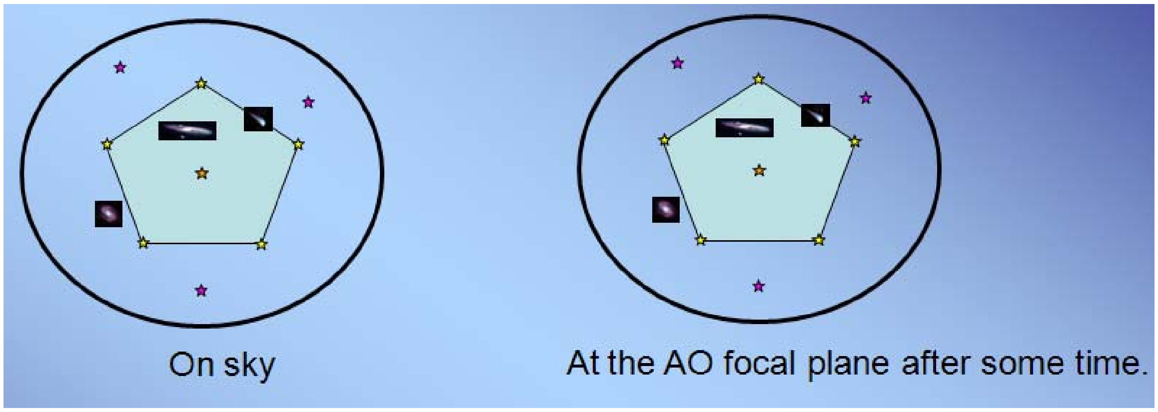


Figure 6 WF scenario – K mirror keeps the stars fixed on the AO focal plane(s) and hence on the WFS channels.

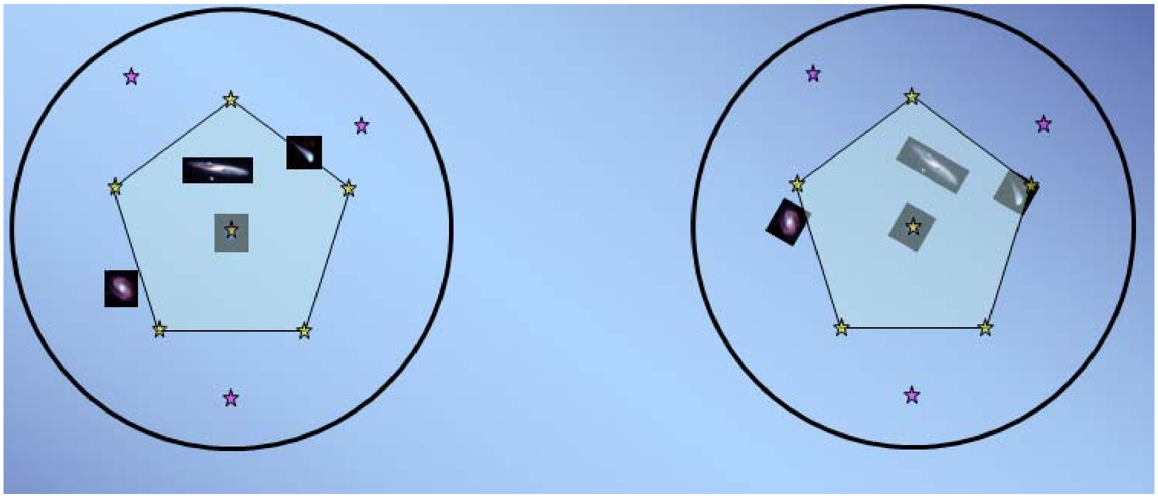


Figure 7 NF scenario in pupil fixed mode (for high contrast imaging) – K mirror keeps pupil fixed. The fact that the sky rotates, does not affect the observations.

Tip-tilt (Focus and astigmatism) sensor:

Low Order Wavefront Sensor (LOWFS) uses a natural star to sense low order modes of the wavefront that are poorly sensed by the multiple LGS. These modes include tip, tilt, focus and astigmatism when the goal is optimizing on axis science performance. When the goal is optimizing science performance averaged over a larger field of view the LOWFS is used to estimate tip, tilt and tilt anisoplanatism modes. A more complete discussion of this issue is included in KAON 492. This also accounts for motion of the Na-layer.

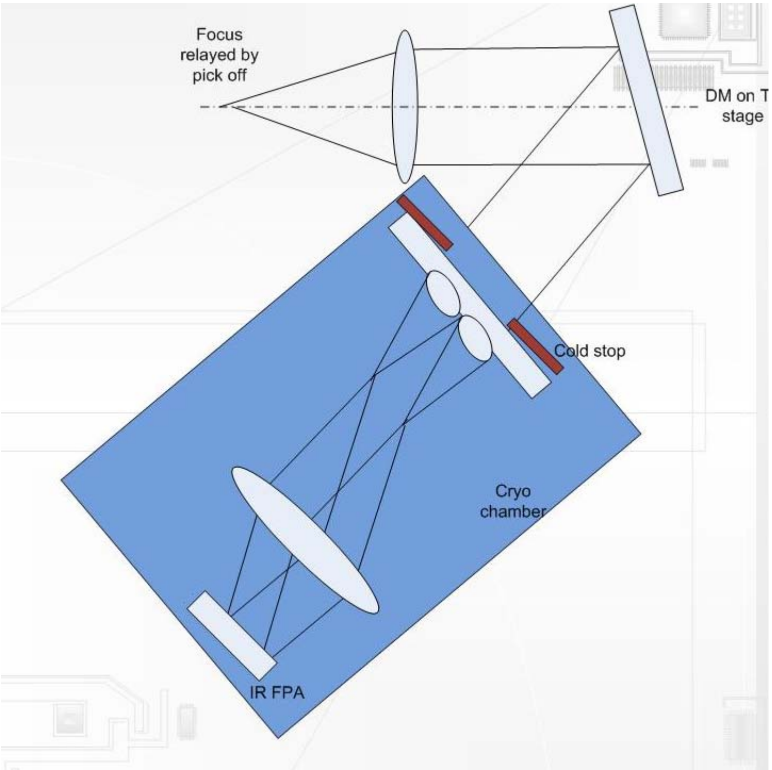


Figure 8 Schematic of the TTFA showing the cold stop and cryogenic part and the front cold end ((there is a window between the cold and the cryogenic chamber).

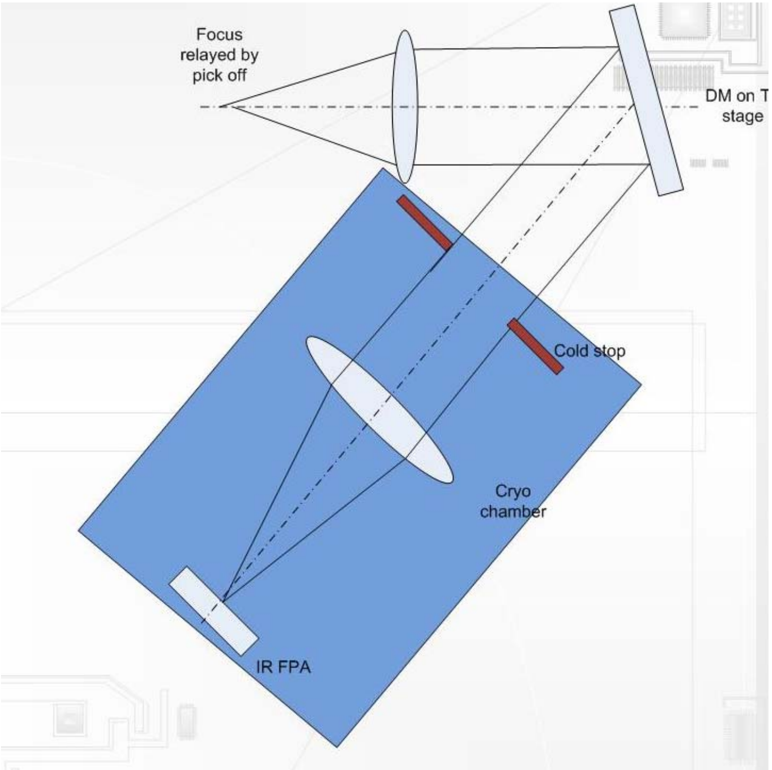


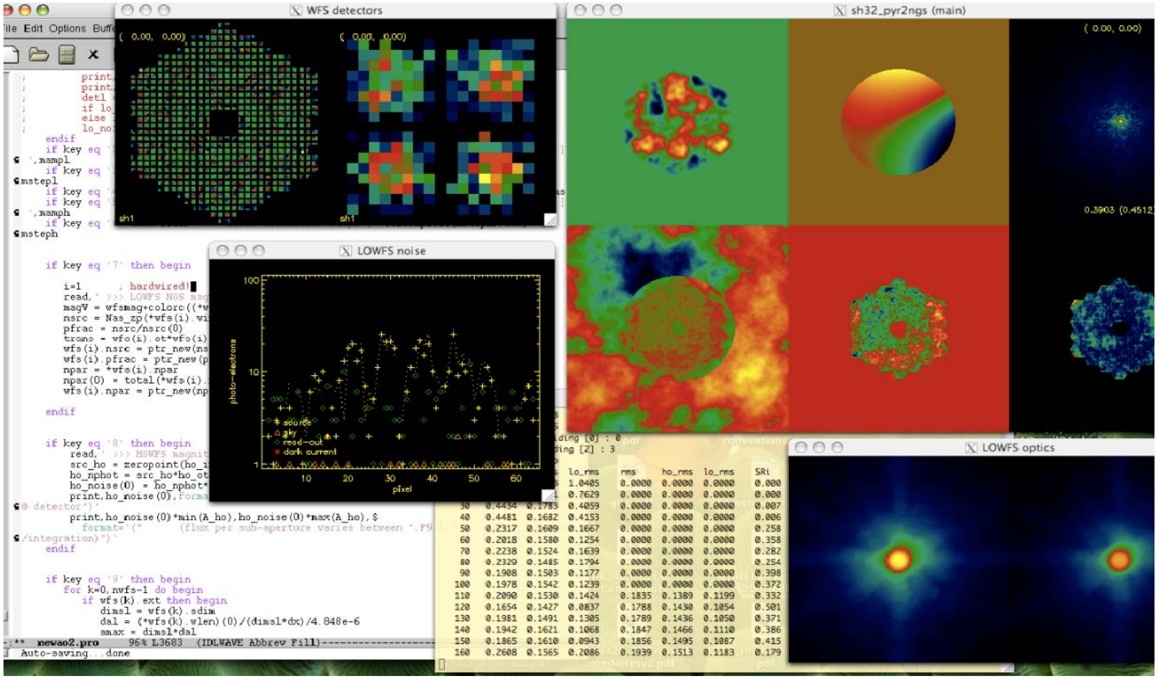
Figure 9 Schematic of the TT sensor design (there is a window between the cold and the cryogenic chamber).

Name	**Number**	**Location**	**Mechanism type**	**Short arm length (mm)**	**Long arm length (mm)**	**Patrol field (arcsec/mm)**	**Arm FoV (arcsec)**	**Acquisition accuracy (mas/μm)**	**Stability*(mas/μm)**	**Dithering implementation**
LGS 1 (corner)	4	LGS WFS unit	θ/φ	65	180	170	5	100	1/0.72	Non required
LGS 2 (center)	4	LGS WFS unit	θ/φ	65	130	170	5	100	1/0.72	Non required
TT star	2	Interim LOWFS unit	θ/φ	50	100	120	5	100	1/0.72	TT/DM
TTFA/TWFS	1	Interim LOWFS unit	θ/φ	50	100	120	5	100	1/0.72	TT/DM
PSF monitor	1	Interim LOWFS unit	θ/φ	50	100	120	5	100	1/0.72	TT/DM

Table 4 Pick off arm specification (subject to change)

"Stability" refers to the movement of a probe arm when the motors are switched off and/or are not moving

There are a total of 2 tip-tilt sensors and one TTFA sensor after the wide-field relay near the d-IFS package. The TT(FA) sensors could use the same pick-off mechanism as the LGS WFS or share pick-offs with d-IFS (if the NGAO instrument team comes up with a different mechanism for d-IFS pick-offs). The TT sensors will use a simple IR H2RG device with selectable pixel pockets. This scheme allows us to deal with extended objects. Some preliminary simulations have been performed to test the validity of the scheme. The TT pick-off mechanism work on the same principle as the LGS pick offs but the lengths are slightly different.



```
;         print.
;         print.
;         detl
;         if lo
;         else
;         lo_no
      endif
      if key eq '
$ ',mampl
      if key eq '
$mstepl
      if key eq '
      if key eq '
$ ',mamph
      if key eq '
$msteph

      if key eq '7' then begin

         i=1         ; hardwired!
         read,' >>> LOWFS NGS mag
         magV = wfsmag+colorc((*w
         nsrc = Nas_zp(*wfs(i).vi
         pfrac = nsrc/nsrc(0)
         trons = wfo(i).ot*wfo(i)
         wfs(i).nsrc = ptr_new(ns
         wfs(i).pfrac = ptr_new(p
         npar = *wfs(i).npar
         npar(0) = total(*wfs(i).
         wfs(i).npar = ptr_new(np

      endif

      if key eq '8' then begin
          read,' >>> HOWFS magnit
          src_ho = zeropoint(ho_i
          ho_nphot = src_ho*ho_ot
          ho_noise(0) = ho_nphot*
          print,ho_noise(0),forma
$@ detector")'
          print,ho_noise(0)*min(A_ho),ho_noise(0)*max(A_ho),$
            format='("      (flux per sub-aperture varies between ",F9
$/integration)")'
      endif

      if key eq '9' then begin
         for k=0,nwfs-1 do begin
            if wfs(k).ext then begin
               dims1 = wfs(k).sdim
               dal = (*wfs(k).wlen)(0)/(dims1*dx)/4.848e-6
               amax = dims1*dal
-:**  newao2.pro      96% L3683   (IDLWAVE Abbrev Fill)------------------
 Auto-saving...done
```

```
ding [0] : 0
ding [2] : 3
                      lo_rms   rms     ho_rms  lo_rms   SRi
                      1.0405   0.0000  0.0000  0.0000   0.000
                      0.7629   0.0000  0.0000  0.0000   0.000
 30   0.4434  0.1783  0.4059   0.0000  0.0000  0.0000   0.007
 40   0.4481  0.1682  0.4153   0.0000  0.0000  0.0000   0.006
 50   0.2317  0.1609  0.1667   0.0000  0.0000  0.0000   0.258
 60   0.2018  0.1580  0.1254   0.0000  0.0000  0.0000   0.358
 70   0.2238  0.1524  0.1639   0.0000  0.0000  0.0000   0.282
 80   0.2329  0.1485  0.1794   0.0000  0.0000  0.0000   0.254
 90   0.1908  0.1503  0.1177   0.0000  0.0000  0.0000   0.398
100   0.1978  0.1542  0.1239   0.0000  0.0000  0.0000   0.372
110   0.2090  0.1530  0.1424   0.1835  0.1389  0.1199   0.332
120   0.1654  0.1427  0.0837   0.1788  0.1430  0.1054   0.501
130   0.1981  0.1491  0.1305   0.1789  0.1436  0.1050   0.371
140   0.1942  0.1621  0.1068   0.1847  0.1466  0.1110   0.386
150   0.1865  0.1610  0.0943   0.1856  0.1495  0.1087   0.415
160   0.2608  0.1565  0.2086   0.1939  0.1513  0.1183   0.179
```

Figure 10 Sample simulation of a TT sensor with 200 mas size object and 30 mas/ pixel scale.

Truth sensors:

Truth or Calibration Wavefront Sensor is used to calibrate biases that arise when using LGS in an adaptive optics system. The biases are principally caused by the elongated nature of the LGS when views by subapertures of the laser guide star wavefront sensor and the changing sodium layer density pro_le. The truth WFS measures this bias by sensing the wavefront from a natural star. These biases are slowly varying and are of a low spatial order. As such, a natural guide stars WFS using long exposures and only measuring the lowest spatial wavefront error is sufficient.

The truth sensor has been estimated to be a 5x5 sensor with detailed error budgets for the same still in the works. The baseline design will have 2 truth sensors as these sensors are to be placed as close as possible to the science instrument. The WF truth sensor will be a IR sensor while the NF field truth sensor as conceptualized to be a visible sensor. With the available information we estimate the plate scale to be 0.65”/pixel for the visible sensor and 200 mas/pixel for the IR TWFS. The visible plate scale was chosen based on the fact that the Strehls are 11% at 5” separation from the on-axis science target and 2% at a off-axis radius of 15”. The median seeing limited spot size is 0.65”.

Since the 20x20 woofer mirror won’t give much correction over the large field, the IR TWFS can either be enabled with a separate MOAO relay to guide on a faint star or split (5%) light from the TTFA sensor. The later scheme is attractive since the TWFS is a slow sensor and needs little light, at the same time it saves on the cost of the AO relay and the real-time computation hardware/ software to correct in yet another direction. The later scheme also helps acquire stars faster than the counterpart. The plate scale is chosen to be the same as the TT sensors. The Strehls are like 20% in H-band and 10% in J-band, well within the range to be able to sense TT on the core of the PSF.

Status of the truth sensor error budget – The EB tool is being revamped to include the TWFS in the control loop. Currently, a new spreadsheet to evaluate truth sensor SNR and centroiding error has been made. In the visible arm one can go down to 21.5 mag with 10 sec. integration time and achieve 35 nm of total TWFS error.

5. NGS sensor: The NGS sensor is to be positioned in the narrow field and will be equipped to pick off stars using a field steering mirrors/ dichroics 40”. The dichroic for splitting light the NGS sensor will be the same as the dichroic for splitting light for the TWFS. Since the TWFS and the NGS sensor are never used simultaneously, it is easiest to use the same articulated dichoric as the first field steering mirror. The NGS and TWFS can share the 1st of the pick off assembly. This one may need to be articulated. There are 2 different cameras to serve as the TWFS and the NGSWFS. I would not recommend using the NGS WFS as the TWFS for the following reasons:
   1. NGS sensors are built for speed (low RON but significant dark current). Even though they are at -20 deg. C ambient, this may be an issue. We don't have exact DC #s for the CCID56 detectors.
   2. At least the CCD50 (e2v) are built to run only at speeds greater than 25 Hz (if I remember correctly). May be the CCID56s can run at 1-0.1 Hz and work at 2000 Hz without a problem, but its possible that this is hard to do.
   3. Just the mechanical complexity of adding another pupil scale with an odd # of sub aps makes me shy away from one NGSand TWFS sensor.

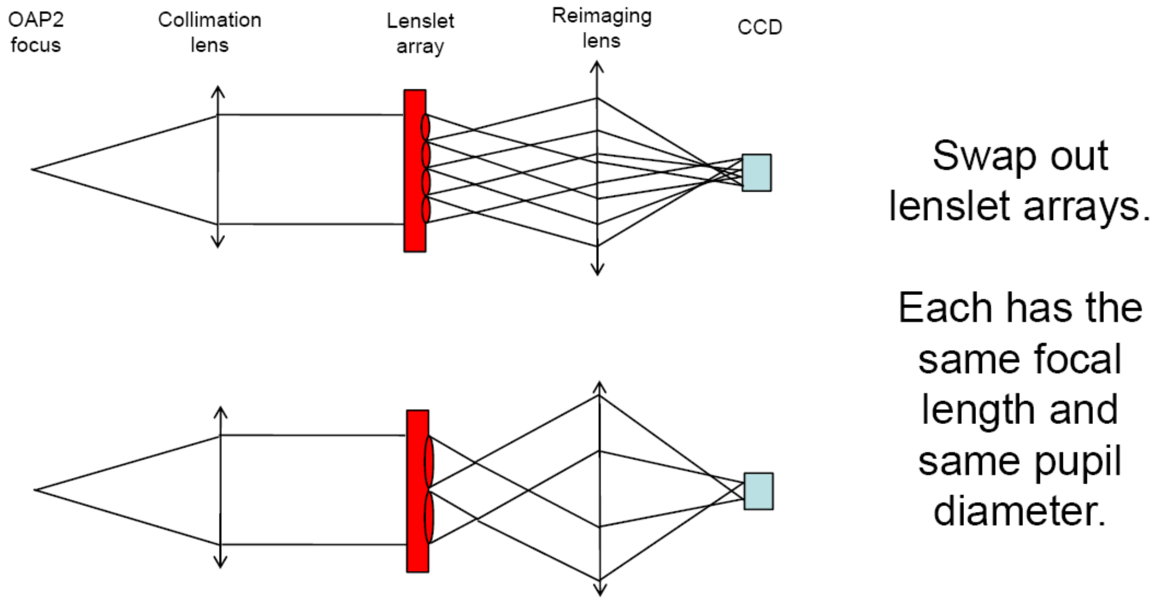


Figure 11 Shows the basic schematic for the lenslet switching scheme for the simplest mechanism to facilitate multiple pupil sampling scales.

WFS type	Length in Z [mm]	Breadth X [mm]	Width Y [mm]	# of sensors	Pick off mech.	Cooling
LGS WFS	200	150	150	9	LGS theta-phi	3 stage Peltier
NGS WFS	450	100	150	1	Chris and Reni's dichroic and steering mirror mech.	3 stage Peltier
NF truth sensor	450	100	150	1	Chris and Reni's dichroic and steering mirror mech.	3 stage Peltier
WF truth sensor	200	150	150	1	TT/ Truth theta-phi	Cryotiger
TT	650	150	150	2	TT/ Truth theta-phi	Cryotiger
TTFA	650	150	150	1	TT/ Truth theta-phi	Cryotiger

Table 5: WFS volume table (subject to change)

NGAO WFS parts	Charecteristics	units	Comments
CCD detectors	128x128 pixels, read out @ 2000 Hz with 3 e- RON and <500 e-/pix/sec DN @ operating temp.	10	9 LGS WFS + 1 NGS WFS
Truth sensor (visible)	10x10 active pixels read out at 0.01-200 Hz with 0.001 e-/pix/sec DN and 3 e- RON	1	truth sensor NF
Truth sensor (IR)	10x10 active pixels read out at 0.01-200 Hz with 0.001 e-/pix/sec DN and 3 e- RON	1	truth sensor WF
IR detectors	Must be able to read in PoI mode. RON < 7 e- with 16 Fowler samples, data rate = 500 Hz at this RON spec for the entire chip, DC = 0.001 e-/pix/sec @ 73K. # of active pixels = 128 x128?	3	2 TT and 1 TTFA sensors
Stages	See table 3	see table 3	separate table
Pick off arms	See Anna's KAON	4 + 4 LGS pick offs + 3 TT pick offs	All LGS pick offs are not the same
Optics	See table 5		

Table 6: List of the major components required for the WFS subsystem

WFS optics table (preliminary)	**Optic**	**Quantity**
NGS WFS	Collimator	1
	Lenslet array (64x64, 32 x 32)	2
	Relay lens	1
	Focusing lens	1
	Field stop (adjustable?)	1
LGS WFS	Collimator	9
	Lenslet array (64x64, 32 x 32, 16x16)	27
	Relay lens	9
	Focusing lens	9
	Field stop (adjustable?)	9
TTFA sensor	Collimator	1
	MEMS mirror	1
	TT stage *(on which MEMS mirror is mounted)*	1
	lenslet array (2x2)	1
	Relay lens	1
	Focusing lens	1
	Cryo chamber	1
	Cold pupil stop	1
TT sensor	Collimator	2
	MEMS mirror	2
	TT stage *(on which MEMS mirror is mounted)*	2
	Relay lens	2
	Focusing lens	2
	Cryo chamber	2
	Cold pupil stop	2
truth sensor (WF)	Collimator	1
	TT stage *(on which MEMS mirror is mounted)*	1
	lenslet array (5x5)	1
	Relay lens	1
	Focusing lens	1
	Cryo chamber	1
	Cold pupil stop	1

truth sensor (NF)	Collimator	1
	Lenslet array (5x5)	1
	Relay lens	1
	Focusing lens	1

Table 7: Table of optics required for the different WFS channels

For more information about pick offs see A. Moore's KAON.

Glossary:

Invisible modes Invisible modes: Some part of the modal content of atmospheric turbulence falls into a category called invisible modes, signifying that they are registered as zero measurements by all the WFSs, but in between the WFS beacons they are non-zero and contribute to the wavefront error seen by a science instrument. In the case of laser guide star (LGS) beacons, these modes also become non-zero for beams focused at infinity in the direction of the LGSs. These modes are by necessity three-dimensional, and are rendered invisible to the AO system by a conspiracy in which modes at different altitudes cancel out within the beam print to leave only piston in each beam, which is not sensed.

Null-modes (also Blind-modes) Null-modes (also Blind-modes): A special case of invisible modes are the so-called null-modes, which are particular to a LGS-based tomography system (as opposed to NGS-based). The null-modes arise from combinations of modes that only cancel out partially and leave, in addition to piston, also linear (i.e. tip/tilt) terms in the beam. When these terms are filtered in a LGS system, due to the tilt determination problem with LGSs, the result is a loss of information that renders the system blind to these modes as well. These differential tilt modes (sometimes called plate-scale modes because of their effect, or null-modes because they belong to the null space of the LGS interaction matrix) are produced by a combination of quadratic modes (e.g. focus, astigmatism) occurring at different altitudes. Hence, in the Kolmogorov model of atmospheric turbulence, these modes have relatively large weights in the turbulence power spectrum, and the impact of not correcting them can be severe

DM Tilt Anisoplanatism Modes The DM Tilt Anisoplanatism Modes are (typically) three modes of deformable mirror commands in an MCAO system that compensate for tilt anisoplanatism over an extended FoV without introducing higher-order wavefront errors. The real time control system controls these modes based upon tip, tilt, focus, and astigmatism measurements from natural guide star (NGS) wavefront sensors (WFS), and possible the focus and astigmatism modes of laser guide star (LGS) WFS measurements.

Object selection Mechanism is a contraption that helps the Multi-object Deployable Near-IR IFU and TT(FA) select field stars/ objects.